

Infrared spectroscopy of nearby radio active elliptical galaxies

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Abstract: In preparation for a study of their circumnuclear gas we have surveyed 60% of a complete sample of elliptical galaxies within 75 Mpc that are radiosources. Some 20% of our nuclear spectra have infrared emission lines, mostly Paschen lines, Brackett γ and [FeII]. We consider the influence of radio power and black hole mass in relation to the spectra. Access to the spectra is provided here as a community resource.

Keywords: infrared: general — active galactic nuclei – galaxies: elliptical – radiosources

1. Introduction

Astrophysicists have striven to learn the detailed physics of Active Galactic Nuclei (AGN) since black holes were first posited as the energy source (Lynden-Bell & Rees 1971). Landmarks were the dusty torus model (Mayes, Pearce & Evans 1985) and the unified scheme of AGN which realised the implications of this geometry (Antonucci *et al.* 1987) from the observer’s standpoint. With the advent of Laser Guide Star Adaptive Optics on 10 metre class telescopes, the opportunity now exists to image the circumnuclear gas in nearby active galaxies and, with the aid of detailed modelling, to understand the physics of the central 10-20 light years of these powerful engines. The purpose of the project our small team has recently commenced is to study a complete sample of galaxies that are radiosources within 40 Mpc. Our work focuses first on the elliptical galaxies, as these are the simplest in terms of stellar population and gas and dust content. We emphasize

nearby galaxies, as at 20 Mpc the diffraction limited resolution of the OSIRIS instrument of the Keck Observatory is 3.2 pc, which is appropriate for disk gas heated by AGN¹.

Central black holes are now believed to be basic constituents of most, if not all, massive galaxies (Magorrian *et al.* 1998, Kormendy 2004). That black hole growth is linked with galaxy formation was realised because of the correlation of BH mass with the properties of the host galaxy (Kormendy 1993; Kormendy & Richstone 1995; Magorrian *et al.* 1998; Gebhardt *et al.* 2000; Ferrarese & Merritt 2000; Barth, Greene & Ho 2005; Greene & Ho 2006), although alternative explanations have been put forth (Jahnke & Maccio 2010).

The deviation of the galaxy mass function from the strict power-law describing dark matter halo masses provides an additional and independent argument in favor of feedback (e.g. Hopkins *et al.* 2008). Feedback is also relevant to the cooling flow problem, particularly with respect to radio AGN (e.g. Bower *et al.* 2006, Best+06). Studies have shown that $\sim 30\%$ of the most massive galaxies are radio continuum sources (e.g., Fabbiano *et al.* 1989; Sadler *et al.* 1989). As a direct manifestation of accretion, and therefore black hole growth, AGN and the consequences of their energy feedback have figured prominently in most current ideas of structure formation (e.g., Granato *et al.* 2004; Springel, Di Matteo & Hernquist 2005; Hopkins *et al.* 2006). Awareness of the importance of black holes has engendered broad interest in the study of the AGN phenomenon itself.

This paper focuses on nuclear activity in nearby galaxies. By selection, most of the objects occupy the faintest end of the AGN luminosity function and have low accretion rates. Although energetically less impressive, low luminosity AGNs deserve scrutiny. By virtue of proposed reasons such as duty cycle (Greene and Ho 2007) or high/low accretions states (see e.g. Hardcastle, Evans & Croston 2006), most AGNs are found in a low luminosity state. Brown *et al.* (2011) argue that all massive early-type galaxies are radiosources, mostly powered by AGN (Slee *et al.* 1994; Ho 1999). The advent of new telescopes and new analysis techniques yield fresh insights into this problem.

The aims of our project are (1) to image nearby ellipticals that are radiosources in 1–2 μm emission lines, (2) to deduce the kinematics of the gas in these lines, (3) to model the density and temperature of the gas, (4) to deduce the mass of the nuclear black hole and the star formation rate in the central 10 pc, (5) to measure gas inflow and outflow rates in non-equilibrium regions of the central structure and (6) summarize these results over the 40 Mpc complete sample of ellipticals, noting correlations with mass, radio power, star formation rate, emission line strength and x-ray power.

In this paper we present the first half of a survey of infrared emission lines carried out at Palomar Observatory and Kitt Peak National Observatory. A number of the galaxies show lines

¹Our sample mean is actually 40 Mpc, where the OSIRIS resolution is 6.4 pc.

of [FeII] which will be followed up with 8 meter class telescopes to accomplish the first goal of our project. Circumnuclear star formation rates are presented for those with Brackett γ in emission.

2. Observations

Brown *et al.* (2011) have compiled a complete sample to $K = 9$ of 396 elliptical galaxies with declination $> -40^\circ$. The 231 detected radiosources² among them are the basic sample for this paper. For diffraction limited imaging of the centers of these galaxies with the Keck and Gemini Observatories we have begun to observe the subsample with 1.4 GHz radio power $> 10^{21}$ W Hz⁻¹ and declination $> -36^\circ$ with the instrumentation listed in Table 1. Sample near infrared spectra are shown in Figure 1–4, in which no redshift correction has been applied. Currently, 20% of the galaxies observed are showing emission lines of [FeII] and/or Br γ , suitable for followup with Keck OSIRIS and Gemini North NIFS integral field unit (IFU) spectrographs³. Our 4 runs (Table 1) with infrared slit spectrographs have netted nuclear spectra of 136 galaxies from the basic sample. This is sufficient to understand the statistics of circumnuclear properties in current epoch ellipticals. Areas of notable incompleteness are the greater Virgo cluster area and the southern hemisphere. Black hole masses in Table 4 are based on the K-band magnitude, following Graham (2007).

At both telescopes all galaxies were observed in two slit positions, A & B, separated by $20''$, in an ABBA sequence of 4–5 minute exposures. Following the object sequence, an 8th magnitude A0V star was observed within an airmass of 0.1 of the object for telluric correction. This was also used for spectrophotometric calibration which employed Vega model fluxes for zeroth visual magnitude. For wavelength calibration an argon lamp was used at Kitt Peak, telluric OH lines at Palomar.

The galaxies with emission lines in these near infrared spectra are listed in Table 2; those without are in Table 3, which also gives the signal to noise ratio of the Kitt Peak spectra at 1.01μ and Palomar spectra at 1.04μ .

3. Data Reduction

The data sets from each observatory required different data reduction packages. The Kitt Peak National Observatory data was reduced using IRAF and the reduction of the Palomar Observatory data was done using Spextool v4.0 beta. (Cushing, Vacca & Rayner 2004).

The first step in the reduction of the Kitt Peak data was the creation of the flat field and dark

²Galaxies for which upper limits were given in Brown’s data table were omitted

³For all but the faintest nuclei, on-axis guiding is possible.

images. The flat field was created by subtracting the flat lamps-off image from the flat lamps-on image. The first and second nights’ data had separate flat field images. The third night’s data set did not contain flat field images, so the flat field image created for night 2 was also used for night 3 objects.

For the standard stars, from the observation files the dark image was subtracted and then the flat field image was divided out. The spectrum was then extracted using the task ‘apall’ in IRAF. As the observation files for the galaxies were taken as ABBA, the galaxy spectra were extracted by creating images for A-B and B-A files which were then divided by the flat field image. The flat fielded A-B and B-A images were then rotated and cropped before running ‘apall’ on each image.

The axes for the spectra were initially in terms of pixels and had to be wavelength calibrated using an argon arc spectrum. This was achieved by picking five prominent peaks in the arc spectrum and identifying the pixel value corresponding to the wavelength of each peak. These five values were then used to interpolate along the entire spectrum.

To create the flat field image used with the Palomar data, IRAF was first used to form a set of flat field images by subtracting the individual flat lamp-off images from the flat lamp-on images. The flat field images created with IRAF were then used in the program ‘xspextool’ in Spextool to create the master flat field. The wavelength calibration was created in ‘xspextool’ using any four galaxy images. The spectra for the galaxies and standard stars were extracted in ‘xspextool’ for A-B and B-A image sets using the master flat field image and the wavelength calibration.

The individual extracted spectra for each galaxy and standard star were then combined using ‘xcombspec’ by selecting all images of each object and scaling all orders with respect to the spectrum with the greatest magnitude in the order 3 spectrum.

Telluric correction was carried out on the combined spectra for each galaxy using the combined standard star spectra with a similar air mass to the galaxy using the Spextool program ‘xtellcor’. The hydrogen absorption lines in the standard star spectrum were rescaled for all orders and the telluric spectrum was constructed. The four orders (3,4,5,6) of the telluric corrected galaxy spectra were merged using ‘xmergeorders’ to give a single spectrum.

Spectra are provided for readers desiring access to them⁴. Cesetti *et al.* (2009) show spectra of NGC 4649, NGC 4697, IC 4296, NGC 4564, NGC 5077, and NGC 5576. Riffel *et al.* (2006) show a spectrum of NGC 2110.

⁴<https://sites.google.com/site/radioactivegalaxies/>

4. Discussion

We compare the radio power of our emission line galaxies with that of the full 397 galaxy Brown *et al.* sample in Figure 6. Infrared emission line nuclei have 6 times stronger radio emission than average at a given K band magnitude (or stellar mass). This does not stand out in the figure, but is a 4σ effect, based on the line in Figure 6 and the scatter in the relation. Both radio and infrared quantities here are from the data table of Brown *et al.* . The whole-sample correlation may be expected on general scaling grounds. A larger galaxy hosts a larger black hole which in turn is responsible for more radio power. Indeed, Ho (2008) shows a strong correlation between radio loudness and x-ray luminosity divided by Eddington ratio. However, if our infrared emission line galaxies stand out, it may be because the BH in these cases is currently being fed by the circumnuclear disk.

We also examined the thermal infrared fluxes of the emission line galaxies. Only UGC 3426 and NGC 3516 had mid-infrared excesses in the AKARI Point Source Catalog (Murakami *et al.* 2007): the former has a $9\mu/2.2\mu$ flux ratio of 1.7 and the latter 1.0. The infrared flux of UGC 3426 rises a further factor of 6 at 18μ . In general, the IRAS fluxes of our near-infrared emission line galaxies do not correlate well with the SFR in Table 2 (Figure 7).

The infrared emission line spectra in Table 2 are not all clones of one another. We compare the Br γ profiles of NGC 2787 and IC 630 in Figure 8. The former has a broad rounded profile, FWHM 750 km/sec. The latter has a sharp profile, FWHM 200 km/sec. The former is a LINER (Balzano 1983), the latter a starburst nucleus in the Veron-Cetty & Veron (2006) catalog. Both are morphologically classified S0. These differences are a strong motivation to explore the spatial structure of the emitting gas at 10 meter class diffraction limited resolution.

In Figure 9 we show the relation between BH mass and SFR. The latter is calculated from Br γ (Kennicutt 1998). SFR uncertainties were calculated from the Br γ uncertainties in Table 2. We draw the uncertainty in the BH mass from Table 4 where both the (K, M_{BH}) and (σ , M_{BH}) relations are employed and the galaxies have secondary distance indicators. It is largely a scatter plot, although the sample is modest, and there is a hint of a correlation in the galaxy scaling sense. If the BH formation process were steady in solar masses/year, one might expect a correlation. If the process is sporadic, dependent for example on fuelling events or circumnuclear disk instability, a correlation would easily be masked. Monitoring Br γ over time or high resolution analysis of any disk could shed light on this. Photoionization by hard radiation from close to the BH could also strengthen Br γ over the value due to star formation.

5. Conclusions

Some 20% of ellipticals in the Brown *et al.* sample have emission line nuclei which will be followed up with IFU spectrographs on AO equipped telescopes. The complete sample may have a

somewhat lower emission fraction, as galaxies with optical emission lines (where this was known) tended to be observed first in order to provide immediate targets for IFU observing. Emission line nuclei are accompanied by more powerful radio sources at a given galaxy mass.

The SFR in these nuclei has a median value of $0.4 \text{ M}_{\odot}/\text{yr}$ and must be intermittent to attain the median black hole mass of $5 \times 10^8 \text{ M}_{\odot}$ of these galaxies.

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Table 1: Runs

#	Telescope	Spectrograph	Dates	Resolution	Slit	Wavelength range
1	Hale 5m Palomar	Triplespec	14–15/9/2011	2600	1'' 4 pixels	1.0–2.4 μm
2	Mayall 4m KPNO	Flamingos	6–11/11/2011	1000	1'' 3 pixels	0.9–1.8 μm
3	Hale 5m Palomar	Triplespec	2/1/2012	2600	1''	1.0–2.4 μm
4	Mayall 4m KPNO	Flamingos	4–9/02/2012	1000	1''	0.9–1.8 μm
5	Hale 5m Palomar	Triplespec	29–30/5/2012	2600	1''	1.0–2.4 μm

Note (1) The following radio active early type galaxies have Palomar Triplespec spectra from Woo *et al.* (2012):

NGC 3245, 3607, 3608, 4261, 4291, 4374, 4459, 4473, 4486, 4564, 4596, 4649, 4697, 4742, 6251, 7052.

None of these galaxies have emission lines in H band.

Note (2) The Triplespec detector format is 2048×1024 ; the Flamingos detector is 2048^2 pixels.

Table 2: Emission line nuclei

NAME	RA 2000	DEC	Run	Br γ Flux ergs/cm ² /s	SFR M _⊙ /yr	Remarks	P _{1.4} W/Hz	redshift
NGC128	00:29:15	+02:51:51	1	1.11±0.04E-15	0.38	1.74 μ unident.	6.0E+20	0.014
NGC524	01:24:48	+09:32:19	1	1.85±0.06E-15	0.18	very broad 1.6 μ	2.1E+20	0.008
NGC547	01:26:01	−01:20:42	1	3.97±0.04E-16	0.18		4.0E+24	0.018
NGC665	01:44:56	+10:25:23	1	6.92±0.03E-16	0.38		6.3E+21	0.018
NGC708	01:52:46	+36:09:06	1	9.1±0.9E-16	0.37		3.5E+22	0.016
NGC1052	02:41:05	−08:15:21	1			1.26, 1.64 μ [FeII]	5.1E+22	0.005
NGC1316	03:22:42	−37:12:29	3	2.81±0.05E-14	1.0	1.26, 1.64 μ [FeII]	7.6E+24	0.006
NGC1521	04:08:19	−21:03:07	3	4.62±0.05E-15	1.5		2.3E+21	0.014
NGC1550	04:19:38	+02:24:35	1	1.27±0.03E-15	0.31		5.3E+21	0.012
NGC2110	05:52:11	−07:27:22	1				3.9E+22	0.008
NGC2128	06:04:34	+57:37:40	1	1.7±0.1E-15	0.31		5.7E+21	0.010
NGC2273	06:50:08	+60:50:45	4			[FeII] lines	6.8E+21	0.006
UGC3426	06:15:36	+71:02:15	1	1.2±0.02E-13	37.2		4.1E+23	0.0135
NGC2768	09:11:38	+60:02:14	1				8.7E+20	0.004
NGC2787	09:19:18	+69:12:12	1	1.63±0.05E-14	0.31		7.3E+19	0.002
NGC2859	09:24:19	+34:30:49	1	7.7±0.4E-15	0.49		7.0E+19	0.006
NGC3065	10:01:55	+72:10:13	1	1.8±0.1E-15	0.18		3.6E+20	0.007
IC630	10:38:33	−07:10:14	3	1.99±0.01E-14	2.1		7.0E+21	0.007
NGC3516	11:06:47	+72:34:07	3			Broad P β	4.8E+21	0.009
NGC4111	12:07:03.1	+43:03:57	5	1.5±0.1E-15	0.04	+	2.5E+20	0.003
NGC5273	13:42:08.3	+35:39:15	5			Broad P γ	1.1E+20	0.004
NGC5322	13:49:15.3	+60:11:26	5	1.9±0.8E-15	0.18		9.1E+21	0.006
NGC7052	21:18:33	+26:26:49	1	2.49±0.07E-15	1.03		7.9E+22	0.016
NGC7550	23:15:16	+18:57:42	1	1.77±0.04E-15	0.86		2.6E+22	0.017
NGC7618	23:19:47	+42:51:10	1	1.94±0.04E-15	0.99		2.3E+22	0.017
NGC7626	23:20:43	+08:13:01	1	1.86±0.06E-15	0.43		2.0E+23	0.011
NGC7743	23:44:21	+09:56:03	1	1.90±0.07E-15	0.09		3.7E+20	0.006

Note (1) the remark on broad Paschen lines applies to all 3 of NGC 3516, 4111 & 5273.

Note (2) the redshifts in column (10) are literature, not observed, values.

Table 3: Absorption line nuclei

NAME	Run	S/N	NAME	Run	S/N
NGC16	1	93	NGC3091	4	3
NGC50	2	8	NGC3100	3	
NGC57	4	14	NGC3158	2	23
NGC315	1	50	NGC3226	2	22
NGC383	1		NGC3245	2	24
NGC410	1	22	NGC3348	4	8
NGC439	3		NGC7454	2	10
NGC474	1	49	NGC3414	2	30
NGC499	4	12	NGC3528	4	20
NGC533	2	2	NGC3607	4	38
NGC596	4	20	NGC3619	4	26
NGC680	4	13	NGC3626	4	14
NGC741	2	7	NGC3665	4	34
NGC883	2	7	NGC3801	4	7
NGC936	2	10	NGC3872	4	17
NGC1016	2	2	NGC3894	4	20
NGC1060	2	15	NGC3957	4	11
NGC1128	1	70	NGC3962	4	31
NGC1167	1	52	NGC3986	4	14
NGC1200	2		NGC4008	4	5
NGC1209	4	10	NGC4036	4	31
NGC1326	3	74	NGC4125	4	31
NGC1399	3	95	NGC4138	4	28
NGC1400	2		NGC4220	4	17
NGC1407	2	8	NGC4589	4	22
NGC1453	1	89	NGC4636	4	25
NGC1531	3	15	NGC4778	4	8
IC359	2		NGC4782	4	17
UGC3024	2	6	NGC4802	4	15
NGC1587	1	64	NGC4825	4	19
NGC1600	2	8	NGC4874	4	8
NGC1573	2	10	NGC4984	4	28
NGC1653	1	70	NGC5044	4	6
NGC1684	1	68	NGC5077	4	16
NGC1726	2	10	NGC5198	4	20
NGC2089	2	8	NGC5444	4	23
NGC2208	2	4	NGC5490	4	21

NGC2256	4	4	NGC5532	4	9
NGC2258	2	28	NGC5631	4	21
NGC2320	4	15	NGC5838	4	16
NGC2314	4	9	IC 459	3	26
NGC2340	2	8	NGC5846	1	32
NGC2418	4	20	NGC6251	1	53
NGC2493	3	85	NGC6482	1	49
NGC2513	2	30	NGC6703	1	67
NGC2612	4	24	NGC6869	2	8
UGC3789	4	15	NGC6903	2	8
NGC2629	3	71	NGC7242	2	7
NGC2685	1	47	NGC7265	2	13
NGC2693	3	82	NGC7391	2	11
NGC2749	2	23	NGC7436	2	22
NGC2872	3	78	NGC7562	2	7
NGC2911	1	19	NGC7600	2	9
NGC2974	4	55	NGC7711	1	83
NGC3078	3	103	NGC7785	1	69
NGC4551	5	50	NGC5796	4	23
NGC5813	5	54	NGC4143	5	87
NGC4203	5	97	NGC4150	5	64
NGC4278	5	83	NGC4710	5	29
NGC4866	5	47	NGC5353	5	93
NGC3557		NGC7426	1		

Table 1. Table 4: BH masses for Brown *et al.* (2011) galaxies

Name	RA	DEC	m-M		K	σ	M(BH) _K	M(BH) _{σ}	Notes
	J2000		mag		2MASS	km s ⁻¹	10 ⁸ M _⊙	10 ⁸ M _⊙	
NGC0016	00:09:04	+27:43:45	33.2	CV	8.78	180	3.18	0.99	
NGC0050	00:14:44	−07:20:42	34.4	CV	8.67	264	11.77	6.40	
NGC0057	00:15:30	+17:19:42	34.4	CV	8.68	326	11.41	17.84	
NGC0128	00:29:15	+02:51:50	33.9	CV	8.52	215	7.90	2.36	
NGC0315	00:57:48	+30:21:29	34.2	CV	7.95	279	18.89	8.37	
NGC0383	01:07:24	+32:24:44	34.2	CV	8.48	279	12.10	8.37	
NGC0410	01:10:58	+33:09:08	34.3	CV	8.38	300	14.41	11.91	
NGC0439	01:13:47	−31:44:49	34.5	CV	8.70	221	12.65	2.70	
NGC0474	01:20:06	+03:24:55	32.5	CV	8.55	163	2.18	0.61	
NGC0499	01:23:11	+33:27:36	33.9	CV	8.73	266	6.94	6.64	
NGC0524	01:24:47	+09:32:19	32.6	CV	7.16	253	8.77	5.20	
NGC0533	01:25:31	+01:45:32	34.4	CV	8.44	279	14.92	8.37	
NGC0547	01:26:00	−01:20:42	34.4	CV	8.49	260	13.85	5.94	
NGC0596	01:32:51	−07:01:53	32.1	CV	7.97	152	2.40	0.44	
NGC0665	01:44:56	+10:25:23	34.4	CV	8.88	190	9.37	1.29	
NGC0680	01:49:47	+21:58:14	33.0	CV	8.73	212	2.70	2.20	
NGC0708	01:52:46	+36:09:07	34.1	CV	8.57	230	9.94	3.27	
NGC0741	01:56:20	+05:37:43	34.4	CV	8.29	291	17.40	10.27	
NGC0883	02:19:05	−06:47:27	34.3	CV	8.89	287	8.84	9.60	
NGC0936	02:27:37	−01:09:21	31.4	CV	6.91	189	3.34	1.26	
NGC1016	02:38:19	+02:07:09	34.1	$D_n - \sigma$	8.58	302	10.02	12.30	
NGC1052	02:41:04	−08:15:20	31.4	SBF	7.45	207	2.16	1.96	
NGC1060	02:43:15	+32:25:29	34.3	$D_n - \sigma$	8.20	303	16.18	12.50	
NGC1128	02:57:42	+06:01:29	34.9	CV	8.98	...	14.72	...	
NGC1167	03:01:42	+35:12:20	34.2	CV	8.64	171	9.68	0.78	
NGC1200	03:03:54	−11:59:29	33.7	CV	8.58	204	6.56	1.83	
NGC1209	03:06:02	−15:36:40	32.8	SBF	8.32	230	3.37	3.27	
NGC1316	03:22:41	−37:12:29	31.7	SBF	5.59	228	16.23	3.14	
NGC1326	03:23:56	−36:27:52	31.4	TF	7.45	118	2.04	0.13	
NGC1399	03:38:29	−35:27:02	31.6	SBF	6.31	346	7.64	23.8	5.1 ± 0.7 (3)
NGC1400	03:39:30	−18:41:17	31.1	SBF	7.81	252	1.06	5.10	
NGC1407	03:40:11	−18:34:48	31.1	SBF	6.70	271	3.12	7.27	

Table 1—Continued

Name	RA	DEC	m-M		K	σ	M(BH) _K	M(BH) _{σ}	Notes
NGC1453	03:46:27	−03:58:08	33.7	$D_n - \sigma$	8.12	328	10.38	18.38	
NGC1521	04:08:18	−21:03:06	34.1	SBF	8.68	242	8.98	4.19	
NGC1531	04:11:59	−32:51:03	30.6	TF	8.42	112	0.36	0.10	
IC0359	04:12:28	+27:42:06	33.8	CV	8.47	...	7.54	...	
NGC1550	04:19:37	+02:24:35	33.6	CV	8.77	308	4.68	13.54	
UGC3024	04:22:26	+27:17:52	34.3	CV	9.25	...	6.10	...	
NGC1587	04:30:39	+00:39:42	32.8	$D_n - \sigma$	8.51	227	2.86	3.07	
NGC1600	04:31:39	−05:05:09	33.5	$D_n - \sigma$	8.04	334	9.41	20.07	
NGC1573	04:35:03	+73:15:44	33.9	$D_n - \sigma$	8.55	303	7.80	12.50	
NGC1653	04:45:47	−02:23:33	33.8	SBF	8.97	246	4.95	4.54	
NGC1684	04:52:31	−03:06:21	33.9	CV	8.69	306	7.25	13.11	
NGC1726	04:59:41	−07:45:19	33.0	$D_n - \sigma$	8.61	246	3.20	4.54	
NGC2089	05:47:51	−17:36:09	33.0	CV	8.81	206	2.57	1.92	
NGC2110	05:52:11	−07:27:22	32.6	CV	8.14	255	3.46	5.41	
NGC2128	06:04:34	+57:37:40	33.1	CV	8.83	261	2.87	6.05	
UGC3426	06:15:36	+71:02:15	33.8	CV	8.97	283	4.65	8.97	
NGC2208	06:22:34	+51:54:34	34.5	CV	9.04	225	9.34	2.94	
NGC2256	06:47:13	+74:14:11	34.3	CV	8.67	221	10.73	2.70	
NGC2258	06:47:46	+74:28:54	33.6	SBF	8.23	287	7.93	9.60	
NGC2273	06:50:08	+60:50:44	32.3	TF	8.48	123	1.77	0.16	
NGC2273	06:50:08	+60:50:44	32.3	TF	8.48	123	1.77	0.16	
NGC2320	07:05:41	+50:34:51	34.5	SN Ia	8.85	315	10.69	15.10	
NGC2314	07:10:32	+75:19:36	34.0	$D_n - \sigma$	8.88	286	6.70	9.44	
IC0459	07:10:38	+50:10:37	34.5	CV	11.29	...	0.99	...	
NGC2340	07:11:10	+50:10:28	34.7	$D_n - \sigma$	8.88	246	13.18	4.54	
UGC3789	07:19:30	+59:21:18	33.3	CV	9.51	...	1.75	...	
NGC2418	07:36:37	+17:53:02	34.2	CV	8.95	247	7.50	4.63	
NGC2493	08:00:23	+39:49:49	33.7	CV	8.83	249	4.94	4.82	
NGC2513	08:02:24	+09:24:47	33.9	$D_n - \sigma$	8.74	274	6.51	7.67	
NGC2612	08:33:50	−13:10:28	31.8	CV	8.78	...	0.83	...	
NGC2629	08:47:15	+72:59:08	33.6	CV	8.85	298	4.37	11.53	
NGC2685	08:55:34	+58:44:03	30.9	TF	8.35	94	0.54	0.04	
NGC2693	08:56:59	+51:20:50	34.0	$D_n - \sigma$	8.60	349	8.27	24.84	

Table 1—Continued

Name	RA	DEC	m-M		K	σ	M(BH) _K	M(BH) _{σ}	Notes
NGC2749	09:05:21	+18:18:47	33.9	$D_n - \sigma$	8.93	260	5.50	5.94	
NGC2768	09:11:37	+60:02:14	31.8	SBF	7.00	181	4.53	1.02	
NGC2787	09:19:18	+69:12:11	29.4	SBF	7.26	194	0.35	1.43	$1.04^{+0.36}_{-0.64}$ (1)
NGC2859	09:24:18	+34:30:48	32.0	TF	8.04	181	2.17	1.02	
NGC2872	09:25:42	+11:25:55	33.9	$D_n - \sigma$	8.72	284	6.75	9.12	
NGC2911	09:33:46	+10:09:07	33.3	CV	8.71	217	3.72	2.47	
NGC2974	09:42:33	−03:41:57	31.7	SBF	6.25	238	8.51	3.87	
NGC3078	09:58:24	−26:55:33	32.7	SBF	7.88	251	4.97	5.01	
NGC3091	10:00:14	−19:38:10	33.6	$D_n - \sigma$	8.09	321	9.20	16.55	
NGC3100	10:00:40	−31:39:51	32.8	CV	8.08	200	4.27	1.66	
NGC3065	10:01:55	+72:10:13	32.5	TF	8.99	160	1.34	0.56	
NGC3158	10:13:50	+38:45:53	34.5	$D_n - \sigma$	8.80	343	11.47	22.84	
NGC3226	10:23:26	+19:53:53	31.9	SBF	8.57	193	1.10	1.40	
NGC3245	10:27:18	+28:30:26	31.6	SBF	7.86	210	1.70	2.10	2.04 ± 0.49 (4)
IC0630	10:38:33	−07:10:14	32.4	CV	8.65	...	1.73	...	
NGC3348	10:47:09	+72:50:22	32.7	$D_n - \sigma$	7.96	238	4.45	3.87	
NGC3414	10:51:16	+27:58:28	32.0	SBF	7.98	237	2.25	3.79	
NGC3516	11:06:47	+72:34:06	33.0	TF	8.51	157	3.34	0.51	
NGC3497	11:07:18	−19:28:17	33.5	CV	8.67	224	5.02	2.88	==NGC3528
NGC3607	11:16:54	+18:03:09	31.8	SBF	6.99	224	4.74	2.88	1.33 ± 0.44 (5)
NGC3619	11:19:21	+57:45:28	32.2	TF	8.57	162	1.57	0.60	
NGC3626	11:20:03	+18:21:26	31.5	SBF	8.16	142	1.17	0.31	
NGC3665	11:24:43	+38:45:46	32.6	TF	7.68	184	5.11	1.11	
NGC3801	11:40:16	+17:43:41	33.4	CV	8.88	198	3.67	1.58	
NGC3872	11:45:49	+13:46:00	33.3	$D_n - \sigma$	8.51	260	4.75	5.94	
NGC3894	11:48:50	+59:24:56	33.2	CV	8.56	265	4.28	6.52	
NGC3957	11:54:01	−19:34:06	32.2	TF	8.69	...	1.36	...	
NGC3962	11:54:40	−13:58:29	32.7	SBF	7.67	233	6.17	3.49	
NGC3986	11:56:44	+32:01:18	33.3	CV	8.98	196	2.89	1.50	
NGC4008	11:58:17	+28:11:32	33.4	$D_n - \sigma$	8.83	226	3.65	3.01	
NGC4036	12:01:26	+61:53:44	31.9	TF	7.56	181	3.18	1.02	
NGC4125	12:08:06	+65:10:26	31.9	SBF	6.86	227	5.95	3.07	
NGC4138	12:09:29	+43:41:06	30.7	SBF	8.20	140	0.51	0.29	

Table 1—Continued

Name	RA	DEC	m-M		K	σ	M(BH) _K	M(BH) _{σ}	Notes
NGC4220	12:16:11	+47:52:59	31.4	TF	8.13	125	1.05	0.17	
NGC4589	12:37:25	+74:11:30	31.7	SBF	7.76	224	2.09	2.88	
NGC4636	12:42:49	+02:41:15	30.8	SBF	6.42	203	3.24	1.78	
NGC4778	12:53:06	−09:12:14	33.9	CV	8.63	232	7.35	3.42	==NGC4759
NGC4782	12:54:35	−12:34:06	33.9	$D_n - \sigma$	7.75	326	16.64	17.84	
NGC4802	12:55:49	−12:03:18	30.3	SBF	8.50	...	0.26	...	
NGC4825	12:57:12	−13:39:52	33.9	CV	8.51	308	8.51	13.54	
NGC4874	12:59:35	+27:57:34	35.0	GCLF	8.86	279	17.49	8.37	
NGC4984	13:08:57	−15:30:58	31.6	TF	7.74	...	1.98	...	
NGC5044	13:15:23	−16:23:04	32.5	SBF	7.71	242	4.56	4.19	
NGC5077	13:19:31	−12:39:21	32.9	$D_n - \sigma$	8.22	256	4.24	5.51	
NGC5198	13:30:11	+46:40:14	33.4	$D_n - \sigma$	8.90	196	3.49	1.50	
NGC5444	14:03:24	+35:07:55	33.7	$D_n - \sigma$	8.84	229	5.13	3.21	
NGC5490	14:09:57	+17:32:43	34.5	SN Ia	8.92	288	10.59	9.77	
NGC5532	14:16:52	+10:48:32	35.1	CV	8.76	294	20.22	10.80	
NGC5631	14:26:33	+56:34:57	32.2	SBF	8.47	171	1.72	0.78	
NGC5796	14:59:23	−16:37:25	32.9	$D_n - \sigma$	8.15	273	4.36	7.53	
NGC5838	15:05:26	+02:05:58	32.3	TF	7.58	266	4.26	6.64	
NGC5846	15:06:29	+01:36:20	32.0	SBF	6.93	239	6.01	3.95	
NGC6251	16:32:31	+82:32:16	35.1	CV	9.03	325	15.63	17.57	6.0 ± 0.2(6)
NGC6482	17:51:48	+23:04:18	33.7	CV	8.37	308	7.82	13.54	
NGC6703	18:47:18	+45:33:01	32.1	SBF	8.25	178	1.95	0.94	
NGC6869	20:00:42	+66:13:39	32.5	SBF	8.71	166	1.83	0.67	
NGC7052	21:18:33	+26:26:48	34.1	CV	8.57	284	9.20	9.12	3.7 ^{+2.6} _{−1.5} (2)
NGC7242	22:15:39	+37:17:55	34.5	CV	8.33	...	17.95	...	
NGC7265	22:22:27	+36:12:35	34.2	CV	8.69	258	9.79	5.72	
NGC7391	22:50:36	−01:32:41	34.0	$D_n - \sigma$	8.63	244	8.08	4.36	
NGC7426	22:56:02	+36:21:40	34.3	CV	8.81	...	9.55	...	
NGC7436	22:57:57	+26:09:00	35.0	CV	9.01	352	15.77	25.90	
NGC7454	23:01:06	+16:23:18	31.9	SBF	8.86	114	0.86	0.11	
NGC7454	23:01:06	+16:23:18	31.9	SBF	8.86	114	0.86	0.11	
NGC7550	23:15:16	+18:57:40	34.2	CV	8.91	255	7.90	5.41	
NGC7562	23:15:57	+06:41:14	33.8	SBF	8.32	256	9.03	5.51	

Table 1—Continued

Name	RA	DEC	m-M		K	σ	M(BH) _K	M(BH) _{σ}	Notes
NGC7600	23:18:53	−07:34:49	33.4	CV	8.91	210	3.56	2.10	
NGC7618	23:19:47	+42:51:09	34.3	CV	9.04	298	7.35	11.53	
NGC7626	23:20:42	+08:13:01	33.7	SBF	8.03	271	10.66	7.27	
NGC7711	23:35:39	+15:18:07	33.8	CV	8.91	180	4.93	0.99	
NGC7743	23:44:21	+09:56:02	31.6	SBF	8.42	84	0.97	0.02	
NGC7785	23:55:19	+05:54:57	33.8	$D_n - \sigma$	8.45	255	7.98	5.41	

^[1]J2000 coordinates, from the NASA/IPAC Extragalactic Database (NED, <http://ned.ipac.caltech.edu>)

^[2]Distance modulus estimate, using the method identified in column 5, as listed in the NASA/IPAC Extragalactic Database (NED, <http://ned.ipac.caltech.edu>)

^[3]Method used in determining the distance modulus. For uniformity, preference was given to Surface Brightness Fluctuation (SBF) distances, followed by Type Ia Supernovae (SNIa) distances, distances based on the Tully-Fisher (TF) relation and $D_n - \sigma$ relations, and redshift distances (CV). The latter are calculated from the recessional velocity (as listed in NED) divided by $H_0 = 72 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (Freedman et al. 2000). Specific references can be found in NED.

^[4]Total K-band magnitude derived from fit extrapolation, as listed in the 2MASS database (<http://irsa.ipac.caltech.edu/cgi-bin/Gator/nph-dd>, Jarrett 2000)

^[5]Average value of the stellar velocity dispersion, computed from all available sources following the precepts of Prugniel and Simien (1996), as listed in the HyperLeda database (<http://leda.univ-lyon1.fr/>).

^[6]Black hole mass estimated from the K-band magnitude listed in column 6, following equation 7 of Vika et al. (2012)

^[7]Black hole mass estimated from the stellar velocity dispersion listed in column 7, following equation 20 of Ferrarese & Ford (2005)

^[8]Black hole masses based on dynamical models applied to spatially resolved kinematics

are listed when available. References are: (1) Sarzi et al. (2001); (2) van der Marel & van den Bosch (1998); (3) Gebhardt et al. (2007); (4) Barth et al. (2001); (5) Gultekin et al. (2009); (6) Ferrarese & Ford (1999)

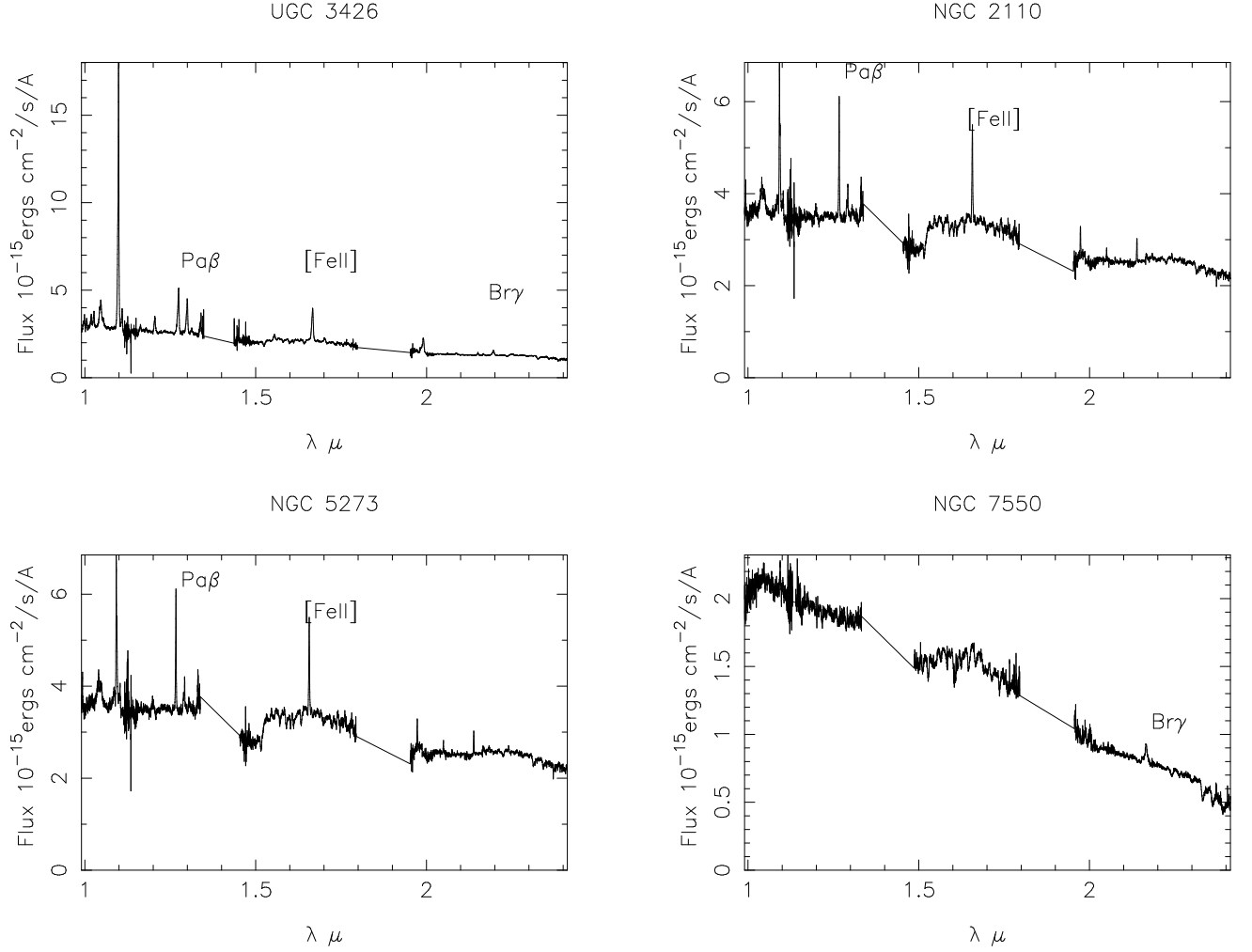


Fig. 1.— Palomar Triplespec spectra ($1''$ aperture) of 3 of the strongest emission line galaxies encountered in the sample so far. [FeII] is the strong line at $1.64\mu m$. Pa β is at 1.28 and Br γ at $2.18\mu m$. NGC 7550 has Br γ in emission, indicating a high star formation rate.

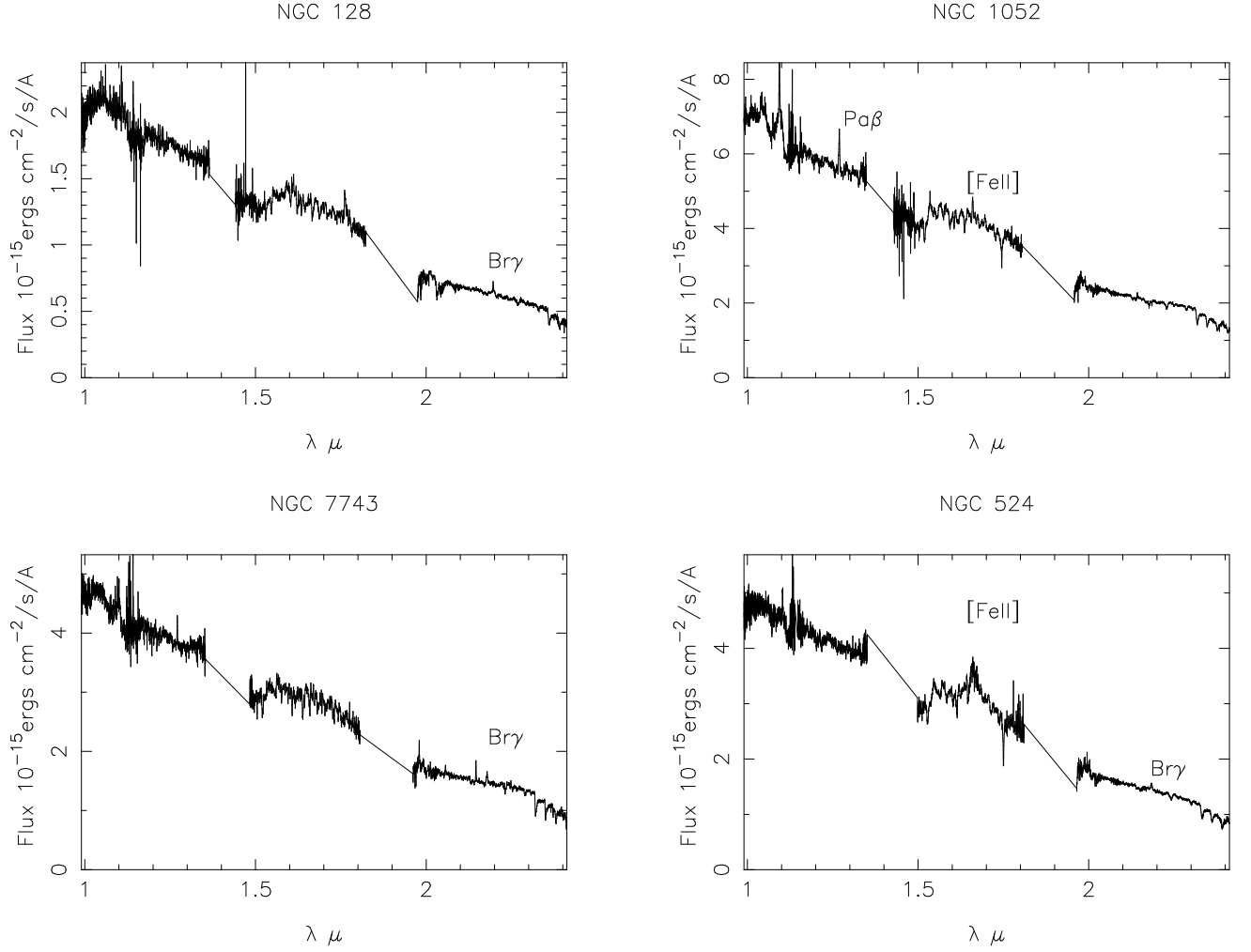


Fig. 2.— Palomar Triplespec spectra of emission line nuclei. The absorption features beyond 2.3μ are CO bandheads from the stellar nucleus.

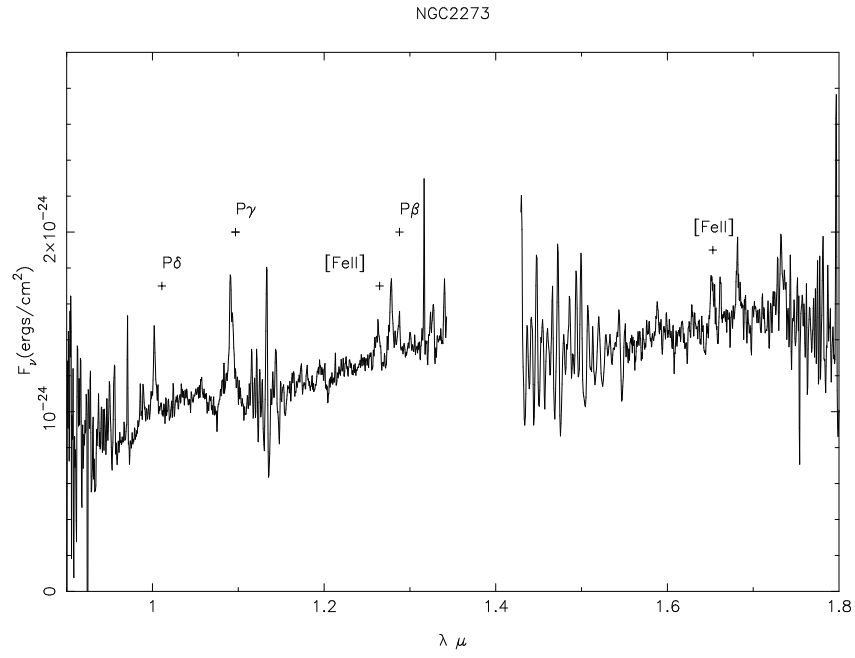


Fig. 3.— Flamingos spectrum of NGC 2273. The Paschen emission lines, such as $P\gamma$ 1.09μ , may thus be exaggerated. Three point boxcar smoothing has been applied.

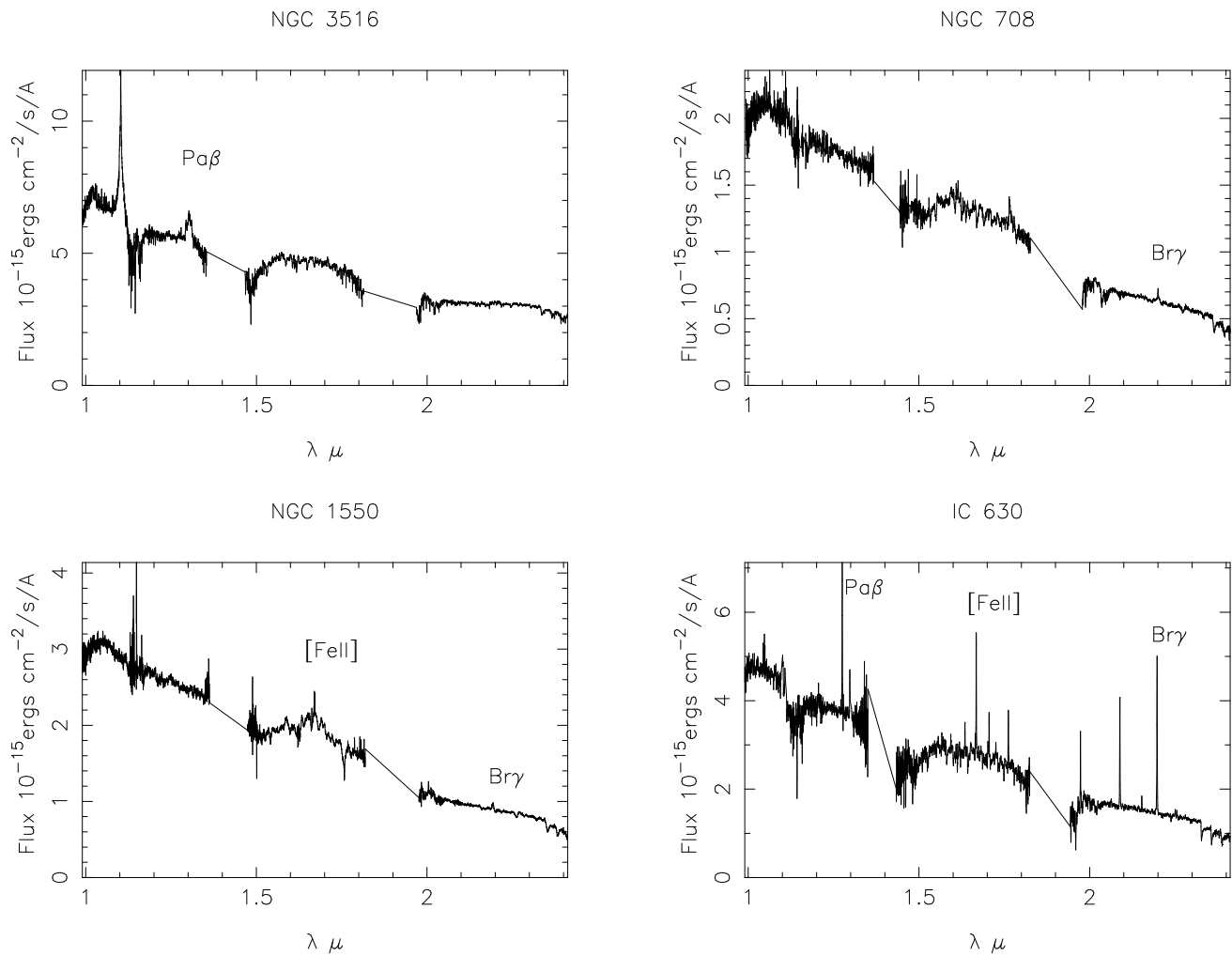


Fig. 4.— Further Palomar spectra of emission line nuclei.

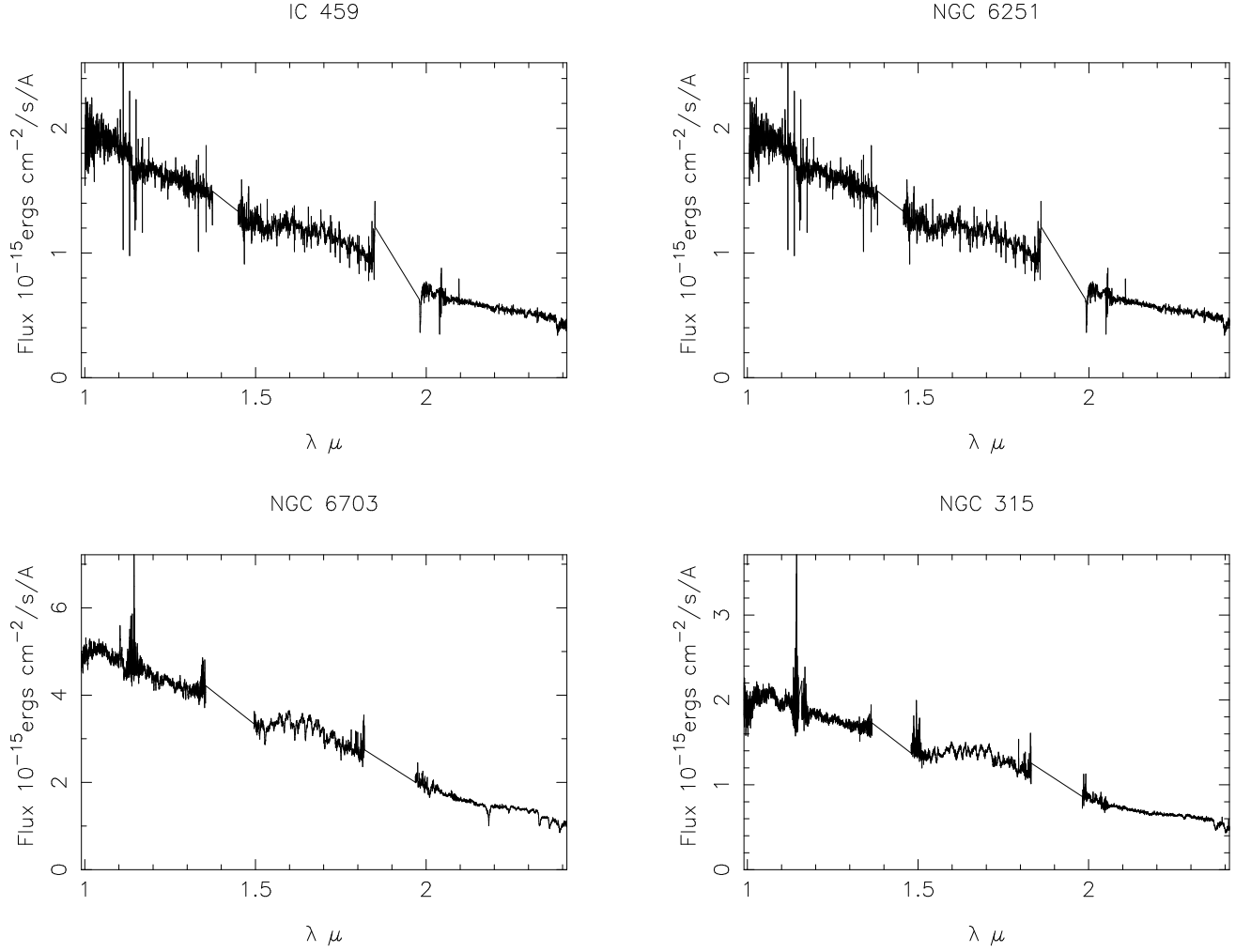


Fig. 5.— Palomar spectra of absorption line nuclei. CO bandheads are seen at the red end of the 2μ window. NG 6703 has $\text{Br}\gamma$ in absorption.

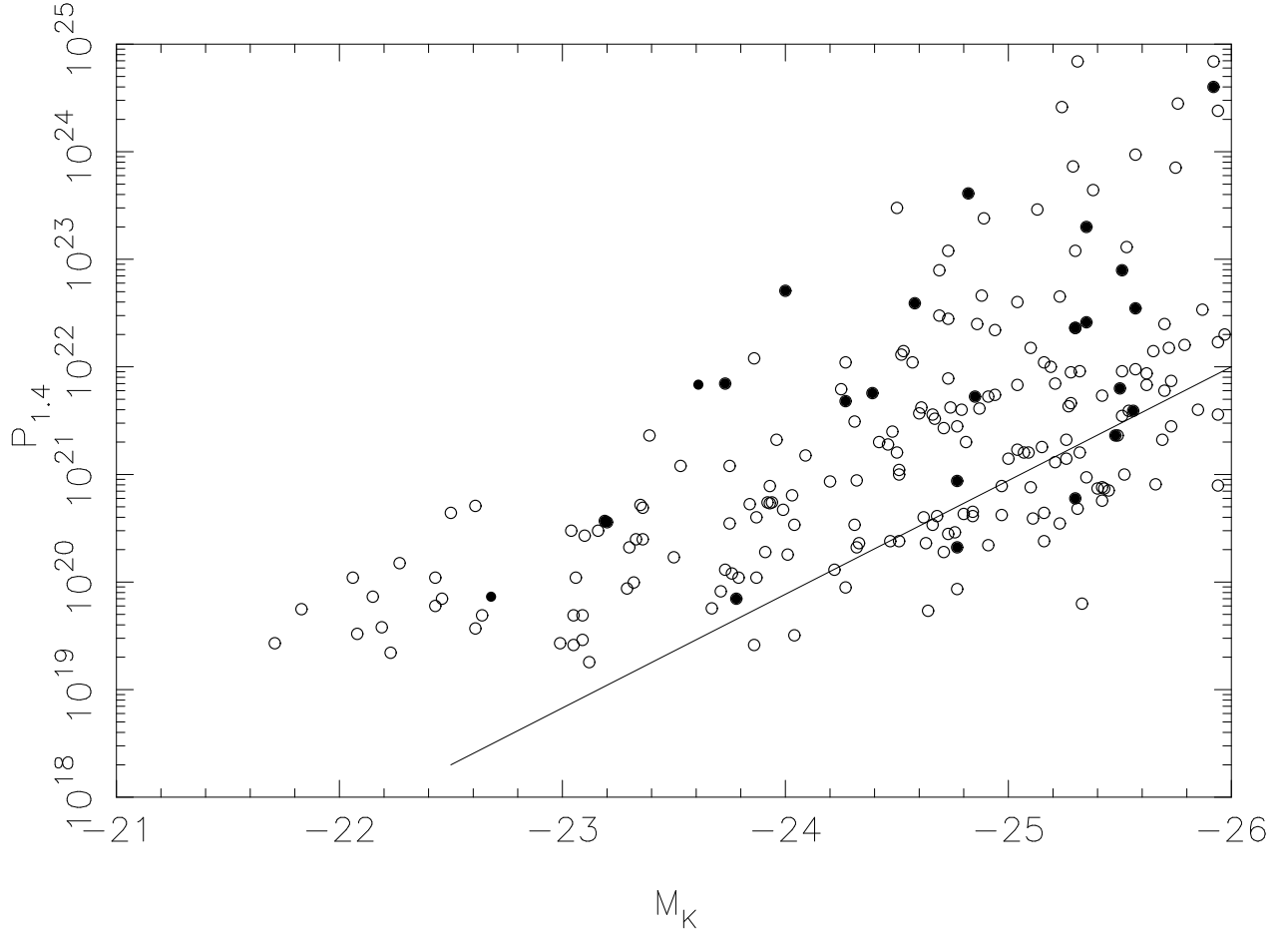


Fig. 6.— Radio power (W/Hz) at 1.4 GHz versus absolute K magnitude, The data are from Brown *et al.* (2011). Galaxies whose radio powers are upper limits are not distinguished. Emission line nuclei are denoted by solid symbols. The line is from the original reference and is the median radio power as a function of K-band luminosity for all ellipticals (including those not detected by NVSS). As this sample only includes radio sources, it is automatically biased towards higher radio powers and is systematically above the line.

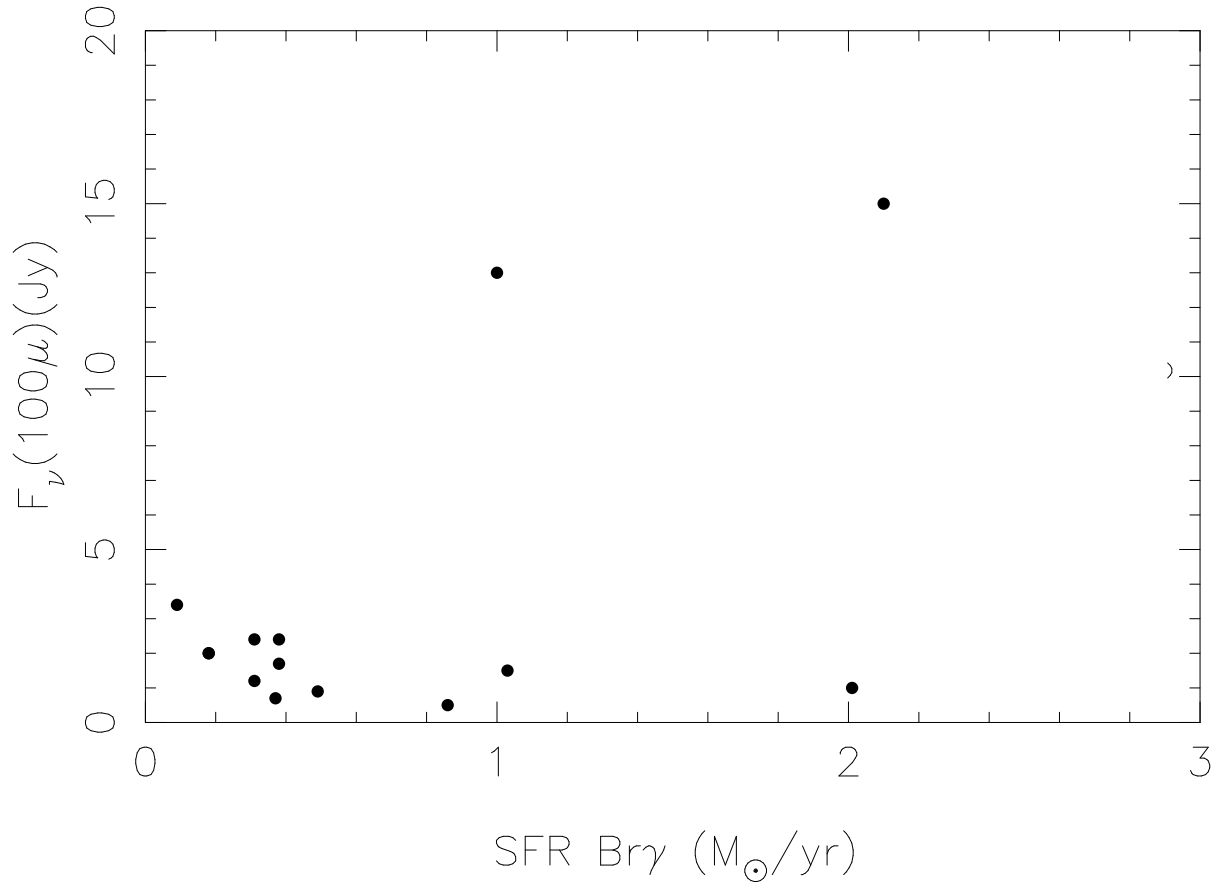


Fig. 7.— Comparison of the Br γ star formation rates with IRAS far infrared fluxes. UGC 3426 is the upper limit at 10 Jy.

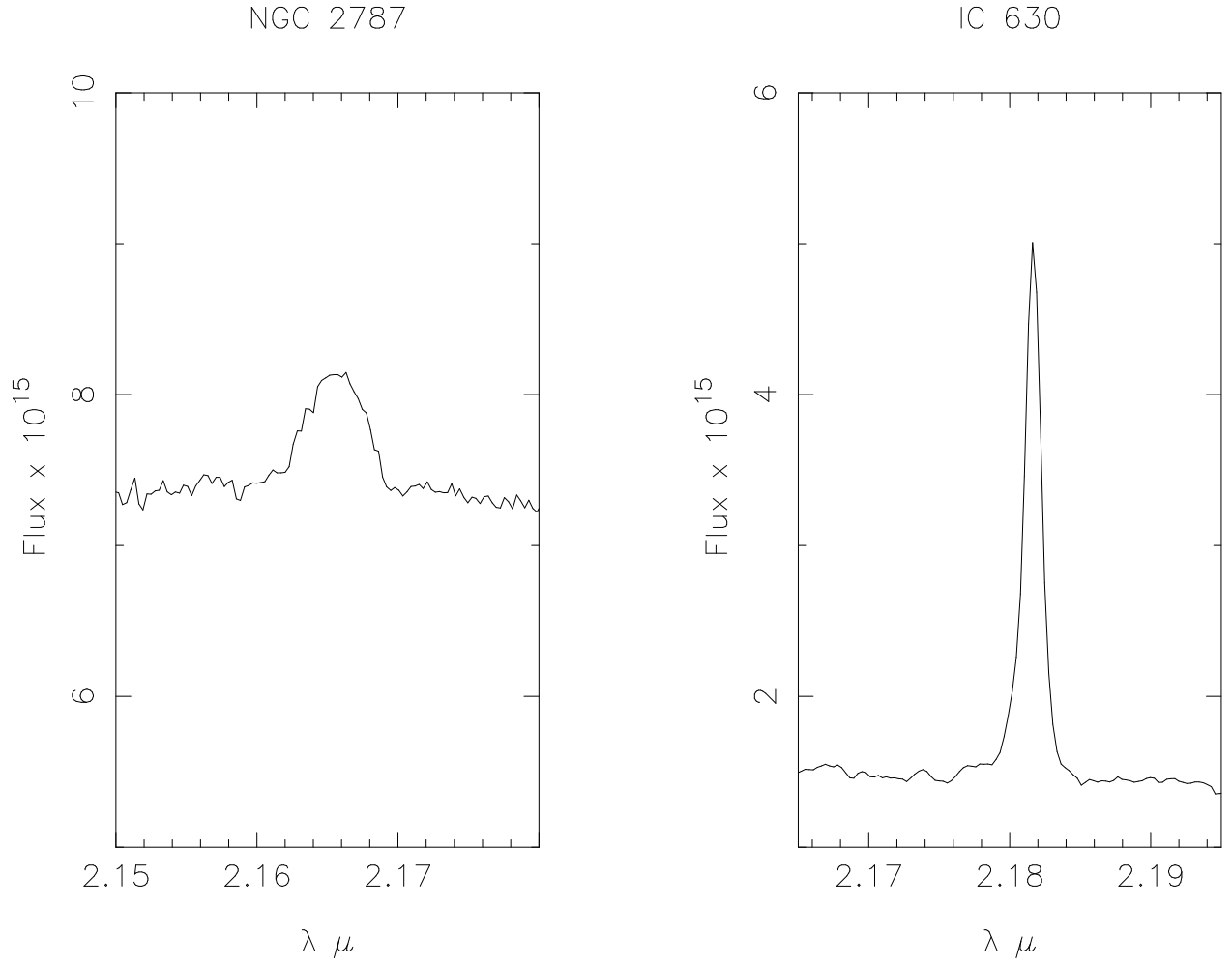


Fig. 8.— Comparison of the Br γ profiles of two galaxies. The spectra have not been corrected to the rest frame.

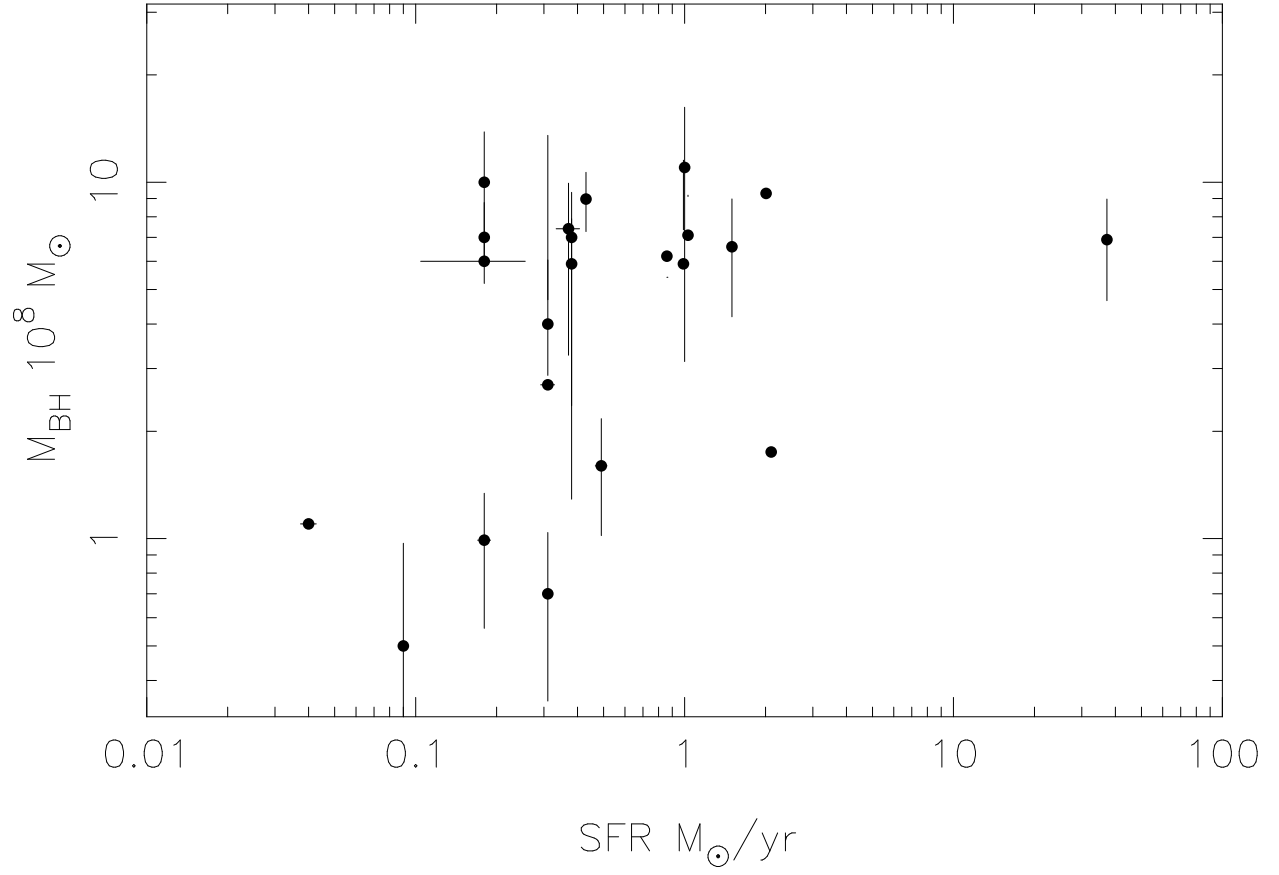


Fig. 9.— SFR and black hole mass from Tables 2 and 4.